



Formation of the BeH⁺ and BeD⁺ Molecules in Be⁺ + H/D Collisions Through Radiative Association

Péter Szabó¹*, Szabolcs Góger² and Magnus Gustafsson³*

¹Department of Physics and Materials Science, University of Luxembourg, Luxembourg, Luxembourg, ²Research Ctr. for Natural Sciences, Institute of Materials and Environmental Chemistry, Hungarian Academy of Sciences, Budapest, Hungary, ³Applied Physics, Division of Materials Science, Department of Engineering Science and Mathematics, Luleå University of Technology, Luleå, Sweden

Cross sections and rate coefficients for the formation of BeH⁺ and BeD⁺ molecules in Be⁺ + H/D collisions through radiative association are calculated using quantum mechanical perturbation theory and Breit-Wigner theory. The local thermodynamic equilibrium limit of the molecule formation is also studied, since the process is also relevant in environments with high-density and/or strong radiation fields. The obtained rate coefficients may facilitate the kinetic modelling of BeH⁺/BeD⁺ production in astrochemical environments as well as the corrosion chemistry of thermonuclear fusion reactors.

OPEN ACCESS

Edited by:

Boutheïna Kerkeni, Manouba University, Tunisia

Reviewed by:

Philippe Halvick, UMR5255 Institut des Sciences Moléculaires (ISM), France Carlo Cavallotti, Politecnico di Milano, Italy

*Correspondence:

Péter Szabó peter88szabo@gmail.com Magnus Gustafsson magnus.gustafsson@ltu.se

Specialty section:

This article was submitted to Astrochemistry, a section of the journal Frontiers in Astronomy and Space Sciences

> Received: 04 May 2021 Accepted: 04 June 2021 Published: 29 July 2021

Citation:

Szabó P, Góger S and Gustafsson M (2021) Formation of the BeH⁺ and BeD⁺ Molecules in Be⁺ + H/D Collisions Through Radiative Association. Front. Astron. Space Sci. 8:704953. doi: 10.3389/fspas.2021.704953 Keywords: molecule formation, radiative association, cross secitons, rate consant, quantum dynamic

1 INTRODUCTION

Lithium, beryllium and boron are surprisingly rare in the Universe compared to other nuclides of similar mass, having abundances more similar to heavy nuclei in meteorites. The widely accepted explanation for this observation is that these elements cannot be produced in normal stellar nucleosynthesis while also being depleted in hydrogen burning processes, producing helium (Burbidge et al., 1957). These elements can still be formed in spallation reactions of the more abundant elements e.g., carbon, nitrogen and oxygen. Out of these three elements, beryllium is the rarest, with a solar abundance of $\log \epsilon_{Be} = 1.15 \pm 0.20$ (with $\log \epsilon_{H} = 12.00$ by definition) (Chmielewski and Müller, 1975).

Even though (or exactly because) lithium, beryllium, and boron have low abundance, these elements can play major roles in observational cosmology. The primordial lithium problem is a famous unsolved problem in modern physics, which states that the current nucleosynthesis model correctly predicts the abundances of deuterium and helium but fails to account for the observed amount of lithium (Fields, 2011). Beryllium could still be formed in extremely metal-deficient stars, and for this reason, the use of beryllium concentrations as a "big bang calorimeter" has been proposed (Pospelov and Pradler, 2011). The presence of BeH molecule in the Sun was first proposed by Wohl (1971). Later Shanmugavel et al. (2008) gave an evidence for the presence of BeH and its isotopologues by observing their A-X transitions in sunspot umbral spectra.

Another important aspect of the beryllium chemistry is that beryllium is one of the main candidates of coating material for fusion reactors (Federici et al., 2001; Paméla et al., 2007). The main issue with the beryllium coating is that Be erodes easily under plasma exposure by releasing Be and Be⁺ into the plasma, where molecules and molecular ions will be formed. In principle, Be, Be⁺, Be²⁺, BeX, BeX⁺, BeX₂ and BeX⁺₂ species can all be important with X being H, D or T. So in order to aid fusion reactor design, the formation and the reactions of these molecules must be understood.

Because of their importance, the BeH and the BeH⁺ molecules have been extensively studied through theoretical calculations in the last decades. Machado and Ornellas (1991) have shown the possible pathways of the BeH⁺ formation from Be^+ and H. For the electronic states $Be^+(^2S) + H(^2S)$ there are two possible correlations to BeH⁺, the $X^{1}\Sigma^{+}$ being the lowest energy attractive state, and the $a^{3}\Sigma^{+}$ being a higher energy repulsive state, not leading to a product. The Be⁺(²P) state has about 4.0 eV excess energy compared to the Be⁺(²S) state, and the process $Be^+(^2P) + H(^2S)$ has three reactive channels $(A^{1}\Sigma^{+}, b^{3}\Pi \text{ and } B^{1}\Pi \text{ in increasing energy order})$ and one nonreactive channel $c^{3}\Sigma^{+}$. Recent ab initio calculations of the BeH⁺ molecule can be found in Refs. (Koput, 2013; Wan et al., 2014), and for BeH a recent review is available (Dattani, 2015). Collisions with electrons can be important in fusion reactor applications, and the $BeH^+ + e^-$ process has been investigated for example in Refs. (Chakrabarti and Tennyson, 2012; Niyonzima et al., 2018; Iacob et al., 2020; Pop et al., 2020; Pop et al., 2021).

The molecule formation in the collision of the Be^+ and H/D particles is forbidden in a strict sense as the relative translation energy of the reactant atoms cannot be stored in the newly formed vibrational degree of freedom of the diatomic molecule without dissociating back to the atomic components. The product can only be stabilized either by the environment (by thermal collisions with free molecules or by interaction with a solid surface) or by emitting a photon. The last process is called radiative association.

The rate coefficients and cross sections for several diatomic molecules with radiative stabilization mechanism have been calculated in the last 20 years with high-level quantum dynamical and semiclassical methods (Singh and Andreazza, 2000; Barinovs and van Hemert, 2006; Antipov et al., 2009; Franz et al., 2011; Andreazza et al., 2012; Andreazza and de Almeida, 2014; Gustafsson et al., 2014; Nyman et al., 2015; Andreazza et al., 2016; Babb and McLaughlin, 2017; Kathir et al., 2017; Szabó and Gustafsson, 2017; Andreazza et al., 2018; Forrey et al., 2018; Szabó and Gustafsson, 2018; Babb et al., 2019a; Babb et al., 2019b; Gustafsson and Forrey, 2019; Zamecnikova et al., 2019; Zamecnikova et al., 2020; Bai et al., 2021). The formation of BeH⁺ molecule through radiative association has not been studied yet. Nevertheless, it might have a contribution to BeH⁺ formation in dust poor regions of the interstellar medium. Furthermore, it would be desirable to learn its importance under plasma conditions as well, since it appears in the corrosion chemistry of thermonuclear fusion reactors.

The aim of this present work is to investigate the BeH⁺ and BeD⁺ production by radiative association in absence of spinorbit and nonadiabatic couplings using quantum dynamical methods.

2 Theory

The cross section of radiative association (RA) can be obtained by quantum mechanical perturbation theory (Babb and Dalgarno, 1995) which yields the Golden rule-like formula

$$\sigma(E_c) = \frac{64\pi^5 f_{\text{stat}}}{3k_{\text{ini}}^2 (4\pi\epsilon_0)} \sum_{J_{j',v'}} \frac{S_{J,j'}}{\lambda_{J_{j',v'},E_c}^3} |M_{J_{jj',v',E_c}}|^2,$$
(1)

where the sum runs over the initial angular momentum (*J*) and final rotational (*j'*) and vibrational (*v'*) quantum numbers. $k_{\text{ini}} = \sqrt{2 \ \mu E_c} / \hbar$ is the wavenumber, where μ is the reduced mass of the colliding fragments, $S_{J,j'}$ is the so-called Hönl-London factor, λ_{J,j',v',E_c} is the wavelength of the emitted photon, E_c is the collision energy, and f_{stat} is the statistical weight factor determined by the initial atomic and final molecular electronic states. The transition dipole matrix elements are defined as

$$M_{J,j',\nu',E_c} = \langle \varphi_{j',\nu'}^{\text{fin}}(r) | D(r) | \chi_{J,E_c}^{\text{ini}}(r) \rangle,$$
(2)

where D(r) is the dipole moment function, $\chi_{j,E_c}^{\text{ini}}(r)$ is the radial part of the energy normalized scattering wavefunction of the initial state and $\varphi_{j',v'}^{\text{fin}}(r)$ is the radial part of the final rovibrational wavefunction, normalized to unity.

The thermal rate constant can be calculated from the cross section according to collision theory as

$$k(T) = \left(\frac{8}{\pi\mu k_B^3 T^3}\right)^{1/2} \int_0^\infty \sigma(E_c) E_c \mathrm{e}^{-\frac{E_c}{k_B T}} \mathrm{d}E_c,\tag{3}$$

where T is the temperature and k_B is the Boltzmann constant.

Besides quantum mechanical perturbation theory, we use Breit-Wigner theory to calculate the RA cross sections and rate constants, since it provides a more controlled treatment of the resonances than perturbation theory (Bennett et al., 2003). According to Breit-Wigner theory (Breit and Wigner, 1936; Bain and Bardsley, 1972), if the interference effects between the direct and resonance mechanisms may be neglected, then the total RA cross section and the rate coefficient can be decomposed as

$$\sigma(E_c) = \sigma_{\rm dir}(E_c) + \sigma_{\rm res}(E_c), \tag{4}$$

$$k(T) = k_{\rm dir}(T) + k_{\rm res}(T), \tag{5}$$

where $\sigma_{\rm res}(E_c)$ and $k_{\rm res}(T)$ are the resonance contributions and $\sigma_{\rm dir}(E_c)$ and $k_{\rm dir}(T)$ are the direct contributions. Since the direct contribution does not contain resonances by definition, $\sigma_{\rm dir}(E_c)$ can be identified as the baseline of the exact cross section which can be obtained either by perturbation theory (using **Eq. 1**) or it can be estimated by the classical trajectory method (Gustafsson, 2013; Nyman et al., 2015). The resonance contribution to the cross section and the rate coefficient can be obtained as (Nyman et al., 2015)

$$\sigma_{\rm res}\left(E_c\right) = \frac{\pi\hbar^2 f_{\rm stat}}{2\mu E_c} \sum_{j,\nu} \frac{(2J+1)\Gamma_{\nu,j}^{\rm tun}\Gamma_{\nu,j}^{\rm rad}}{\left(E_c - \epsilon_{\nu,j}\right)^2 + \frac{1}{4} \left(\Gamma_{\nu,j}^{\rm tun} + \Gamma_{\nu,j}^{\rm rad}\right)^2},\tag{6}$$

$$k_{\rm res}(T) = \hbar^2 f_{\rm stat} \left(\frac{2\pi}{\mu k_B T}\right)^{3/2} \sum_{j,\nu} \frac{(2J+1)e^{-\frac{\epsilon_{j,\nu}}{k_B T}}}{1/\Gamma_{\nu,j}^{\rm tun} + 1/\Gamma_{\nu,j}^{\rm rad}},\tag{7}$$

where v and j are vibrational and rotational quantum numbers of the quasi-bound states. $\epsilon_{j,v}$ is the energy of the quasi-bound state that corresponds to the peak position of resonances. $\Gamma_{v,j}^{\text{tun}}$ is the width associated with tunneling and $\Gamma_{v,j}^{\text{rad}}$ is the width due to the radiative decay to all lower-lying bound and quasi-bound levels j', ν' , of the final electronic state.

The resonance contribution to the rate constant in **Eq. 7**, according to the conventional Breit-Wigner formula, assumes that the quasi-bound states are populated only by tunneling through the potential barrier. This is appropriate in the low gas density and low radiation field limit (Forrey et al., 2018). If there are other modes of populating the quasi-bound states (interaction with the radiation field or third body collisions), an upper limit for the resonance contribution is obtained by replacing **Eq. 7** by its local thermal equilibrium (LTE) approximation (Gustafsson and Forrey, 2019)

$$k_{\rm res}^{\rm LTE}(T) = \hbar^2 f_{\rm stat} \left(\frac{2\pi}{\mu k_B T}\right)^{3/2} \sum_{j,\nu} (2J+1) \Gamma_{\nu,j}^{\rm rad} e^{-\frac{c_{j,\nu}}{k_B T}},$$
(8)

where the sum runs over the j and v quantum numbers of the quasi-bound states.

3 Computational Details

The electronic structure calculations were carried out with the MOLPRO package (Werner, 2015). The potential energy curve used in the dynamical calculations was calculated with the explicit correlated internally contracted multireference configuration interaction method (icMRCI-F12) with Davidson correction, using the aug-cc-pVQZ-F12 basis set as implemented in MOLPRO. All calculations were carried out in the $C_{2\nu}$ symmetry group. The molecular orbitals were constructed using the state averaged complete active space self-consistent field (CASSCF) method with an active space consisting of four electrons on eight orbitals $(6a_1, 1b_1, 1b_2)$ $1b_{2}$, $0a_{2}$). Five states with A_{1} irreducible representation were included in the state average with equal weights in order to obtain a stable wavefunction at each configuration. For the calculation of the permanent dipole function we employed the full configuration interaction (FCI) method with aug-cc-pVQZ basis set. The dipole was calculated as an expectation value of the dipole operator.

Dynamical calculations of radiative association cross sections were carried out according to the quantum mechanical perturbation theory (PT) and Breit-Wigner (BW) method. In the PT treatment the vibrational Hamiltonian was represented with a DVR basis (Colbert and Miller, 1992) for the calculations of bound states. The Numerov method was used for the calculation of the scattering wavefunctions. The total angular momentum ranged from J = 0 to 60 and the energy range was $E_c = 10^{-5}$ eV–8.0 eV. An energy resolution of $\Delta E = 10-6$ eV was used in the entire energy range so that narrow resonances are reasonably well represented. The resonance parameters, which are needed in the BW method, were computed with the program LEVEL (Le Roy, 2007). The value of the statistical weight factor is $f_{\text{stat}} = 1/4$ for reaction (9) and the non-zero Hönl-London factors are $S_{I \rightarrow J+1} = J + 1$ and $S_{I \rightarrow I-1} = J$ (Hansson and Watson, 2005).

4 RESULTS AND DISCUSSION

There is only one potential energy curve, that of the ground state $({}^{1}\Sigma^{+}$, see **Figure 1**), which can provide an available pathway for the molecule formation in the following reaction



FIGURE 1 | Potential energy (black) and electric dipole curves (green and blue) for reaction (9) used in the dynamical calculations.

$$\operatorname{Be}^{+}({}^{2}S) + \operatorname{H}({}^{2}S) \to \operatorname{BeH}^{+}({}^{1}\Sigma^{+}) + \hbar\omega.$$
(9)

The corresponding permanent dipole moment in the center of mass fixed frame of reference, and the dipole after isotope substitution, are also displayed in **Figure 1**. (See **Appendix A** for details on the relation between the dipole moments of the BeH⁺ and the BeD⁺ molecules). It is clear from **Figure 1** that the dipoles correlate correctly to the free Be⁺ + H/D limit as the atomic distance goes to infinity. The ab initio data of the potential and dipole curves are tabulated in **Appendix B**.

The cross sections for the BeH⁺ molecule calculated with the quantum mechanical perturbation theory are shown in Figure 2A. Due to the attractive potential energy curve and the deep potential well, the efficiency of the RA process is more dominant at lower collision energies and it is still significant up to $E_c = 1.0 \text{ eV}$. Above this energy value, nonreactive processes may dominate in the collisions: bremsstrahlung under 10.0 eV and ionization and electronic excitation of the atoms above (Kramida et al., 2017). The cross section curves show a rich resonance structure, especially above 10^{-3} eV where the very narrow resonance peaks are close to each other. Several of these resonances contain contributions from quasi-bound to quasi-bound transitions. Nevertheless, these transitions may have a significant contribution to the radiative association of colliding fragments (Bennett et al., 2003). This unusual way of molecule formation can be efficient, since the decay into the continuum manifold can be slower than the radiative relaxation to the stable ro-vibrational quantum states. The isotope substitution slightly changes the dynamics of the RA (see Figure 2B): the base-line of the cross section is diminished,



according to a known scaling law (Gustafsson, 2013), and the resonance structure is affected as well.

The PT method has some peculiarities in the description of the resonances. The evaluation of the cross section even for a diatomic molecule can be computationally challenging at high energies. Even with a very fine energy resolution we cannot describe all of the narrow resonances, especially above 0.1 eV (see **Figure 2A**). This means that many of the resonances are missing and others are not perfectly resolved. Owing to this, the corresponding rate constant, computed by numerical integration of (**Eq. 3**), may not be converged. Furthermore, the PT method does not account for radiative broadening of the quasibound states (Le Roy, 2007; Zamecnikova et al., 2020). Therefore, it will give resonances in the cross section where the total width $\Gamma_{v,j}^{tun} + \Gamma_{v,j}^{rad}$ in **Eq. 6** is replaced by $\Gamma_{v,j}^{tun}$. Even if a rate constant was computed from the perfectly resolved PT cross section it would

thus result in the LTE rate constant (Gustafsson and Forrey, 2019).

To bypass the shortcomings of the PT method, we also calculated the cross sections with the Breit-Wigner method (see **Figure 2B**), which makes possible the direct evaluation of the rate constant, **Eq. 7**. In our calculations we used the baseline of the PT cross sections as the direct contribution to the BW method (**Eqs. 4**, 5). **Figure 2A** displays how the BW method provides a comprehensive and accurate description of the resonances in the entire energy range, especially at high energies where the PT method fails to provide resonances.

The rate constants for the formation of BeH⁺ and BeD⁺ molecules are displayed in Figure 3. As expected from the cross sections (Figure 2B), the rate constant of the BeD⁺ is smaller with a factor of 2-3, compared to those of BeH⁺ using all of the methods. Due to the insufficient description of the resonances in the PT method, it underestimates the rate constants in the whole temperature range. The magnitude of the rate constants is comparable to formation by RA of other diatomic molecules (Nyman et al., 2015), but it is still much smaller than rates of common chemical exchange reactions. Based on this, the RA mechanism is likely negligible in a higher density astrochemical environment, where complex chemical networks dominate the molecule production owing to the high diversity and density of chemical species. Nevertheless, RA might be an important reaction channel in the low density, dust-poor regions of the interstellar space, since it is a direct and sometimes the only route to the molecule formation.

Since the production of BeH⁺ and BeD⁺ may be important in a variety of environments, as discussed above, we have also included LTE rate constants in **Figure 3**. Those are larger than the conventional BW rate constant by up to a factor of seven. This may have to be taken into account in studies of astrophysical environments with high density or, in particular, strong radiation fields. It may also play a role in studies of fusion reactors.



FIGURE 3 | Rate constants as a function of temperature for the formation of BeH⁺ and BeD⁺ calculated with PT, BW and LTE methods.

TABLE 1 Fitting parameters of the extended Arrhenius curves for the formation of
BeH ⁺ and BeD ⁺ that are obtained from BW rate coefficients.

BeH ⁺ 5-10 896.8638 1.0808 -3. BeH ⁺ 10-25 1.3073 -0.6473 13. BeH ⁺ 25-100 2.8886 -0.0125 -6. BeH ⁺ 100-500 3.2310 0.1118 -8. BeH ⁺ 500-800 5.8565 -0.2856 186 BeH ⁺ 800-2000 18.6999 -0.8485 677 BeH ⁺ 2000-10,000 77.4333 -1.3379 1682 BeD ⁺ 5-10 0.2575 -0.4381 0.33	K) 0175
BeH+ 10-25 1.3073 -0.6473 13. BeH+ 25-100 2.8886 -0.0125 -6. BeH+ 100-500 3.2310 0.1118 -8. BeH+ 500-800 5.8565 -0.2856 186 BeH+ 800-2000 18.6999 -0.8485 677 BeH+ 2000-10,000 77.4333 -1.3379 1682 BeD+ 5-10 0.2575 -0.4381 0.33	
BeH* 25–100 2.8886 -0.0125 -6. BeH* 100–500 3.2310 0.1118 -8. BeH* 500–800 5.8565 -0.2856 186 BeH* 800–2000 18.6999 -0.8485 677 BeH* 2000–10,000 77.4333 -1.3379 1682 BeD* 5–10 0.2575 -0.4381 0.335	752
BeH+ 100–500 3.2310 0.1118 -8. BeH+ 500–800 5.8565 -0.2856 186 BeH+ 800–2000 18.6999 -0.8485 677 BeH+ 2000–10,000 77.4333 -1.3379 1682 BeD+ 5–10 0.2575 -0.4381 0.3	102
BeH* 500-800 5.8565 -0.2856 186 BeH* 800-2000 18.6999 -0.8485 677 BeH* 2000-10,000 77.4333 -1.3379 1682 BeD* 5-10 0.2575 -0.4381 0.335	1041
BeH* 800-2000 18.6999 -0.8485 677 BeH* 2000-10,000 77.4333 -1.3379 1682 BeD* 5-10 0.2575 -0.4381 0.33	6320
BeH* 2000-10,000 77.4333 -1.3379 1682 BeD* 5-10 0.2575 -0.4381 0.3	8126
BeD ⁺ 5–10 0.2575 –0.4381 0.3	7012
	.1593
RoD ⁺ 10.50 0.0742 0.0800 4	337
DED 10-50 0.9743 0.0099 -4.	2360
BeD ⁺ 50-300 1.1061 0.3783 -24	1499
BeD ⁺ 300-800 2.5625 -0.4374 229	7081
BeD ⁺ 800-2000 6.3406 -0.9020 594	5047
BeD ⁺ 2000-10,000 18.8163 -1.2728 1376	.7414

Following the recommendation of (Antipov et al., 2009), the BW cross section is used to obtain the rate constant for the formation of the BeH^+ and BeD^+ molecules in the interstellar medium. The extended Arrhenius equation

$$k(T) = A\left(\frac{T}{300}\right)^{\alpha} e^{-\beta/T}$$
(10)

is used to fit the BW rate constants. Owing to the complex structure of the reaction rates, the corresponding curves are split into several temperature ranges which are fitted individually. The fitting parameters are summarized for the formation of BeH+ and BeD⁺ in **Table 1** according to different temperature ranges. These rate constants will be made available in the database KIDA (Gustafsson et al., 2012; Wakelam et al., 2012).

5 CONCLUSION

In this work, we studied the radiative association formation of the BeH⁺ and BeD⁺ molecules. The cross sections and the rate constants are calculated by using quantum mechanical perturbation and the Breit-Wigner theory. The obtained data can be utilized in the modelling of the BeH⁺ abundance in the

REFERENCE

- Andreazza, C. M., de Almeida, A. A., and Borin, A. C. (2016). The Radiative Association of P and O Atoms. *Mon. Not. R. Astron. Soc.* 457, 3096–3100. doi:10.1093/mnras/stw116
- Andreazza, C. M., and de Almeida, A. A. (2014). The Formation of AlF by Radiative Association. MNRAS 437, 2932–2935. doi:10.1093/mnras/stt2109
- Andreazza, C. M., de Almeida, A. A., Vichietti, R. M., and Ceccatto, D. T. (2012). Radiative Association of Ti and O Atoms. *Monthly Notices R. Astronomical Soc.* 427, 833–838. doi:10.1111/j.1365-2966.2012.21682.x
- Andreazza, C. M., de Almeida, A. A., and Vichietti, R. M. (2018). Formation of AlCl by Radiative Association. MNRAS 477, 548–551. doi:10.1093/mnras/sty753
- Antipov, S. V., Sjölander, T., Nyman, G., and Gustafsson, M. (2009). Rate Coefficient of CN Formation through Radiative Association: A Theoretical Study of Quantum Effects. J. Chem. Phys. 131, 074302. doi:10.1063/1.3196179

interstellar medium and for its production in thermonuclear fusion reactors. Furthermore, the resonance contribution of the rate constant under local thermal equilibrium conditions is calculated. Such conditions can be particularly important in fusion reactors, as well as in astrochemical environments with strong radiation fields.

DATA AVAILABILITY STATEMENT

The original contributions presented in the study are included in the article/**Supplementary Material**, further inquiries can be directed to the corresponding authors.

AUTHOR CONTRIBUTIONS

PS and MG contributed to conception and design of the study. PS and GS performed the calculations. PS and MG wrote the manuscript. All authors contributed to manuscript, read, and approved the submitted version.

ACKNOWLEDGMENTS

This article is based upon work from COST Action CA18212—Molecular Dynamics in the GAS phase (MD-GAS), supported by COST (European Cooperation in Science and Technology). Support from the Kempe Foundation is gratefully acknowledged. MG acknowledges support from the Knut and Alice Wallenberg Foundation. Computational resources provided by the Swedish National Infrastructure for Computing (SNIC) at HPC2N are acknowledged.

SUPPLEMENTARY MATERIAL

The Supplementary Material for this article can be found online at: https://www.frontiersin.org/articles/10.3389/fspas.2021.704953/full#supplementary-material

- Babb, J. F., and Dalgarno, A. (1995). Radiative Association and Inverse Predissociation of Oxygen Atoms. *Phys. Rev. A*. 51, 3021–3026. doi:10.1103/ physreva.51.3021
- Babb, J. F., and McLaughlin, B. M. (2017). Radiative Association of C(³P) and H⁺: Triplet States. MNRAS 468, 2052–2057. doi:10.1093/mnras/stx630
- Babb, J. F., Smyth, R. T., and McLaughlin, B. M. (2019). Dicarbon Formation in Collisions of Two Carbon Atoms. ApJ 876, 38. doi:10.3847/1538-4357/ab1088
- Babb, J. F., Smyth, R. T., and McLaughlin, B. M. (2019). Radiative Association of Atomic and Ionic Carbon. ApJ 884, 155. doi:10.3847/1538-4357/ab43cb
- Bai, T., Qin, Z., and Liu, L. (2021). Radiative Association for the Formation of MgO. MNRAS 500, 2496.
- Bain, R. A., and Bardsley, J. N. (1972). Shape Resonances in Atom-Atom Collisions. I. Radiative Association. J. Phys. B: Mol. Phys. 5, 277–285. doi:10.1088/0022-3700/5/2/024
- Barinovs, Ğ., and van Hemert, M. C. (2006). CH⁺ Radiative Association. *ApJvol.*636, 923–926. doi:10.1086/498080

- Bennett, O. J., Dickinson, A. S., Leininger, T., and Gadéa, F. X. (2003). Radiative Association in Li⁺H Revisited: the Role of Quasi-Bound States. MNRAS 341, 361–368. doi:10.1046/j.1365-8711.2003.06422.x
- Breit, G., and Wigner, E. (1936). Capture of Slow Neutrons. *Phys. Rev.* 49, 519–531. doi:10.1103/physrev.49.519
- Burbidge, E. M., Burbidge, G. R., Fowler, W. A., and Hoyle, F. (1957). Synthesis of the Elements in Stars. *Rev. Mod. Phys.* 29, 547–650. doi:10.1103/ revmodphys.29.547
- Chakrabarti, K., and Tennyson, J. (2012). Electron Collisions with the BeH⁺ Molecular Ion in the R-Matrix Approach. *Eur. Phys. J. D* 66, 31. doi:10.1140/epjd/e2011-20663-4
- Chmielewski, Y., and Müller, E. A. (1975). The Solar Beryllium Abundance. *Astron. Astrphys.* 42, 37.
- Colbert, D. T., and Miller, W. H. (1992). A Novel Discrete Variable Representation for Quantum Mechanical Reactive Scattering via the S-matrix Kohn Method. J. Chem. Phys. 96, 1982–1991. doi:10.1063/1.462100
- Dattani, N. S. (2015). Beryllium Monohydride (BeH): Where We Are Now, after 86years of Spectroscopy. J. Mol. Spectrosc. 311, 76–83. doi:10.1016/ j.jms.2014.09.005
- Federici, G., Skinner, C. H., Brooks, J. N., Coad, J. P., Grisolia, C., Haasz, A. A., et al. (2001). Plasma-material Interactions in Current Tokamaks and Their Implications for Next Step Fusion Reactors. *Nucl. Fusion* 41, 1967–2137. doi:10.1088/0029-5515/41/12/218
- Fields, B. D. (2011). The Primordial Lithium Problem. Annu. Rev. Nucl. Part. Sci. 61, 47–68. doi:10.1146/annurev-nucl-102010-130445
- Forrey, R. C., Babb, J. F., Stancil, P. C., and McLaughlin, B. M. (2018). Rate Constants for the Formation of CS by Radiative Association. MNRAS 479, 4727–4734. doi:10.1093/mnras/sty1739
- Franz, J., Gustafsson, M., and Nyman, G. (2011). Formation of Carbon Monoxide by Radiative Association: a Quantum Dynamical Study. MNRAS 414, 3547–3550. doi:10.1111/j.1365-2966.2011.18654.x
- Gustafsson, M., Antipov, S. V., Franz, J., and Nyman, G. (2012). Refined Theoretical Study of Radiative Association: Cross Sections and Rate Constants for the Formation of SiN. J. Chem. Phys. 137, 104301. doi:10.1063/1.4750029
- Gustafsson, M. (2013). Classical Calculations of Radiative Association in Absence of Electronic Transitions. J. Chem. Phys. 1387, 074308. doi:10.1063/1.4790862
- Gustafsson, M., and Forrey, R. C. (2019). Semiclassical Methods for Calculating Radiative Association Rate Constants for Different Thermodynamic Conditions: Application to Formation of CO, CN, and SiN. J. Chem. Phys. 150, 224301. doi:10.1063/1.5090587
- Gustafsson, M., Monge-Palacios, M., and Nyman, G. (2014). The Rate Constant for Radiative Association of HF: Comparing Quantum and Classical Dynamics. J. Chem. Phys. 140, 184301. doi:10.1063/1.4874271
- Hansson, A., and Watson, J. K. G. (2005). A Comment on Hönl-London Factors. J. Mol. Spectrosc. 233, 169–173. doi:10.1016/j.jms.2005.06.009
- Iacob, F., Pop, N., Mezei, J., Niyonzima, S., Laporta, V., Chakrabarti, K., et al. (2020). Reactive Collisions between Electrons and BeT⁺. *AIP Conf. Proc.* 2218, 050010.
- Kathir, R. K., Nyman, G., and Gustafsson, M. (2017). The Rate Constant for Formation of HCl through Radiative Association. MNRAS 470, 3068–3070. doi:10.1093/mnras/stx1444
- Koput, J. (2013). Ab Initio ground-state Potential Energy Functions of Beryllium Monohydride Ions: BeH⁺ and BeH⁻. J. Chem. Phys. 139, 104309. doi:10.1063/ 1.4820403
- Kramida, A., Ralchenko, Y., and Reader, J.NIST ASD Team (2017). NIST Atomic Spectra Database (Ver. 5.3). [Online]. Available at: http://physics.nist.gov/asd (National Institute of Standards and Technology, Gaithersburg, MD April 25, 2017).
- Machado, F. B. C., and Ornellas, F. R. (1991). A Theoretical Investigation of the Low-lying Electronic States of the Molecule BeH⁺. J. Chem. Phys. 94, 7237–7244. doi:10.1063/1.460207

- Niyonzima, S., Pop, N., Iacob, F., Larson, Å., Orel, A. E., Mezei, J. Z., et al. (2018). Low-energy Collisions between Electrons and BeD⁺. *Plasma Sourc. Sci. Technol.* 27, 025015. doi:10.1088/1361-6595/aaabef
- Nyman, G., Gustafsson, M., and Antipov, S. V. (2015). Computational Methods to Study the Formation of Small Molecules by Radiative Association. *Int. Rev. Phys. Chem.* 34, 385–428. doi:10.1080/0144235x.2015.1072365
- Paméla, J., Matthews, G. F., Philipps, V., and Kamendje, R. (2007). An Iter-like wall for Jet. J. Nucl. Mater. 363-365, 1–11. doi:10.1016/j.jnucmat.2006.12.056
- Pop, N., Iacob, F., Mezei, J. Z., Niyonzima, S., Motapon, O., and Schneider, I. F. (2020). Dissociative Recombination and Vibrational Excitation of Molecular Cations with Electrons: Application to H2⁺, BeH⁺ and Their Isotopomers. *J. Phys. Conf. Ser.* 1412, 172004. doi:10.1088/1742-6596/1412/17/172004
- Pop, N., Iacob, F., Niyonzima, S., Abdoulanziz, A., Laporta, V., Reiter, D., et al. (2021). Reactive Collisions between Electrons and BeT⁺: Complete Set of thermal Rate Coefficients up to 5000 K. At. Data Nucl. Data Tables 139, 101414. doi:10.1016/j.adt.2021.101414
- Pospelov, M., and Pradler, J. (2011). Primordial Beryllium as a Big Bang Calorimeter. Phys. Rev. Lett. 106, 121305. doi:10.1103/physrevlett.106.121305
- Le Roy, R. J.. University of Waterloo Chemical Physics Research Report CP-663, 2007.
- Shanmugavel, R., Bagare, S. P., Rajamanickam, N., and Kumar, B. (2008). Identification of Beryllium Hydride Isotopomer Lines in sunspot Umbral Spectra. Serbian Astron. J. 176, 51–58. doi:10.2298/saj0876051s
- Singh, P. D., and Andreazza, C. M. (2000). The Formation of CN and CN⁺ by Direct Radiative Association. *ApJ* 537, 261–263. doi:10.1086/309023
- Szabó, P., and Gustafsson, M. (2017). A Surface-Hopping Method for Semiclassical Calculations of Cross Sections for Radiative Association with Electronic Transitions. J. Chem. Phys. 147, 094308. doi:10.1063/1.5000573
- Szabó, P., and Gustafsson, M. (2018). Formation of the NH Molecule and its Isotopologues through Radiative Association. MNRAS 483, 3574. doi:10.1093/ mnras/sty3246
- Wakelam, V., Herbst, E., Loison, J.-C., Smith, I. W. M., Chandrasekaran, V., Pavone, B., et al. (2012). A Kinetic Database for Astrochemistry (KIDA). *ApJS* 199, 21. doi:10.1088/0067-0049/199/1/21
- Wan, M., Wang, F., and Cao, Q. (2014). Low-lying Electronic States of BeH⁺ with the Effect of Inner Electrons. *Mol. Phys.* 112, 2184–2194. doi:10.1080/ 00268976.2014.893037

Werner, H.-J. MOLPRO, Version 2015.1, a Package of Ab Initio Programs, 2015. Wohl, H. (1971). On Molecules in Sunspots. *Sol. Phys.* 16, 362.

Zamecnikova, M., Gustafsson, M., Nyman, G., and Soldan, P. (2020). Formation of CO⁺ by Radiative Association II. *MNRAS* 492, 3794.

Zamecnikova, M., Soldan, P., Gustafsson, M., and Nyman, G. (2019). Formation of CO⁺ by Radiative Association. *MNRAS* 489, 2954.

Conflict of Interest: The authors declare that the research was conducted in the absence of any commercial or financial relationships that could be construed as a potential conflict of interest.

Publisher's Note: All claims expressed in this article are solely those of the authors and do not necessarily represent those of their affiliated organizations, or those of the publisher, the editors and the reviewers. Any product that may be evaluated in this article, or claim that may be made by its manufacturer, is not guaranteed or endorsed by the publisher.

Copyright © 2021 Szabó, Góger and Gustafsson. This is an open-access article distributed under the terms of the Creative Commons Attribution License (CC BY). The use, distribution or reproduction in other forums is permitted, provided the original author(s) and the copyright owner(s) are credited and that the original publication in this journal is cited, in accordance with accepted academic practice. No use, distribution or reproduction is permitted which does not comply with these terms.

APPENDIX A: SCALING OF THE DIPOLE AFTER ISOTOPE SUBSTITUTION

In a center of mass coordinate system the dipole moment (in atomic units) of the ${\rm BeH^+}$ molecule can be defined as

$$D_{\rm BeH} = \frac{m_{\rm Be}q_{\rm H} - m_{\rm H}(1 - q_{\rm H})}{M_{\rm BeH}}r,$$
 (A1)

where $m_{\rm H}$, $m_{\rm Be}$, $M_{\rm BeH}$ are the masses of the H, Be atom and BeH⁺ molecule, respectively. Furthermore, $q_{\rm H}$ is the charge on the hydrogen atom in atomic units and r is the distance between the two atoms in bohrs. In order to obtain the dipole for the isotope substituted molecule (BeD⁺), we define its dipole in the same way

$$D_{\rm BeD} = \frac{m_{\rm Be}q_{\rm H} - m_{\rm D}(1 - q_{\rm H})}{M_{\rm BeD}}r,$$
 (A2)

where M_{BeD} is the mass of the BeD⁺ molecule. The charge on the deuterium atom (q_{H}) has to be the same as in **Eq. A1**. This allows us to calculate D_{BeD} easily, if D_{BeH} is already known. To this end q_{H} can be expressed from **Eq. A1** as

$$q_H = \frac{M_{\rm BeH} D_{\rm BeH} + m_H r}{r M_{\rm BeH}}$$
(A3)

and it can be plugged into Eq. A2 to get the desired scaling of $D_{\rm BeD}$ as a function of $D_{\rm BeH}$.

APPENDIX B: TABULATED AB INITIO DATA

Table A1 contains the ab initio data of the potential energy and electric dipole moment for a Be⁺-H pair, computed as described in **Section 3**. The potential energy curve displayed in **Figure 1** is the interaction potential, i.e. it contains an energy offset compared with the ab initio data in **Table A1**, which ensures that $V(r) \rightarrow 0$ as $r \rightarrow \infty$.

TABLE A1 | Ab initio potential energy, *V*, and electric dipole moment, *D*, of Be⁺–H in the ground state, $X^{1}\Sigma^{+}$, at interatomic separations *r*.

r (Å)	$V(E_h)$	D(ea ₀)
0.45	-13.72760745	-1.34780423
0.50	-13.99631578	-1.32660933
0.55	-14.20621370	-1.31008045
0.60	-14.37115938	-1.29837354
0.65	-14.50117698	-1.29049573
0.70	-14.60377981	-1.28537473
0.75	-14.68472202	-1.28202015
0.80	-14.74844237	-1.27961554
0.85	-14.79839713	-1.27754556
0.90	-14.83731193	-1.27530251
0.95	-14.86734704	-1.27245684
1.00	-14.89022477	-1.26863886
1.05	-14.90732138	-1.26353732
1.10	-14.91974454	-1.25689052
1.15	-14.92838943	-1.24848242
1.20	-14.93398274	-1.23813543
1.25	-14.93711644	-1.22567164
1.30	-14.93827381	-1.19455262
1.35	-14.93785070	-1.17076458
1.40	-14.93616505	-1.14336688
1.45	-14.93386362	-1.11266893
1.50	-14.93048002	-1.07820348
1.55	-14.92646921	-1.04104611
1.60	-14.92203712	-0.99896333
1.70	-14.91241431	-0.90720251
1.80	-14.90236514	-0.80656569
1.90	-14.89244513	-0.70683974
2.00	-14.88300635	-0.60384924
2.10	-14.87425582	-0.50759169
2.20	-14.86632585	-0.42328540
2.30	-14.85925866	-0.35474203
2.50	-14.84769839	-0.27184449
2.70	-14.83931443	-0.25759695
3.00	-14.83117253	-0.32096316
3.20	-14.82791973	-0.39040360
3.30	-14.82675376	-0.42676849
3.50	-14.82508298	-0.49829905
4.00	-14.82300857	-0.65433151
4.50	-14.82233430	-0.78205792
5.00	-14.82206595	-0.89522199
5.50	-14.82193700	-1.00158073
6.00	-14.82186667	-1.10444789
8.00	-14.82175639	-1.50114871
10.00	-14.82172632	-1.88851025
12.00	-14.82171163	-2.27252428
16.00	-14.82169857	-2.21202420